

# The sound of DHOST

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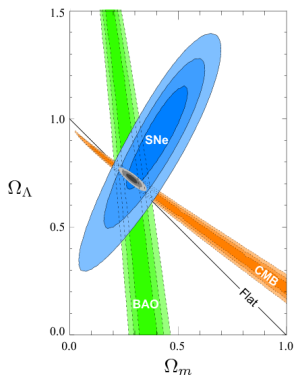
## Summary of the talk

$$\Phi' = \frac{G_{\text{N}} M}{r^2} + \frac{\gamma_1 G_{\text{N}}}{4} M''$$

- 1 DHOST models
- 2 Sound waves
- 3 Constraints on DHOST models

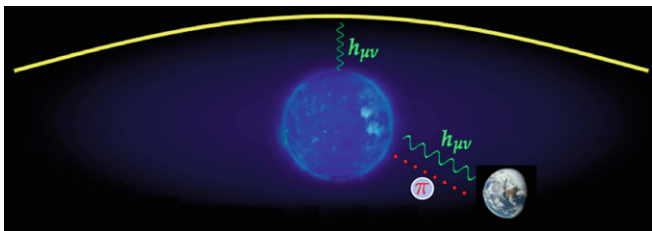
## Modifying gravity: why?

Many reasons; mostly evoked to solve dark energy problems (large cosmological constant problem, coincidence problem)



# Modifying gravity: why with a scalar field?

- Simplest additional degree of freedom
- Many theories related in specific regimes



## Modifying gravity with a scalar field: how?

### DHOST theories

Most complicated way of adding the simplest degree of freedom

# Modifying gravity with a scalar field: how?

DHOST theories

Langlois &amp; Noui '15 (+ others...)

$$S_{\text{DHOST}} = \int d^4x \sqrt{-g} \left[ F_0(\varphi, X) + F_1(\varphi, X) \square\varphi + F_2(\varphi, X) R \right. \\ \left. + \sum_{I=1}^5 A_I(\varphi, X) L_I^{(2)} \right]$$

$$L_1^{(2)} = \varphi^{\mu\nu} \varphi_{\mu\nu}$$

$$L_2^{(2)} = (\square\varphi)^2$$

$$L_3^{(2)} = \square\varphi \varphi^\rho \varphi_{\rho\sigma} \varphi^\sigma \quad \text{with } \varphi_{\mu\nu} = \nabla_\nu \nabla_\mu \varphi, \varphi_\mu = \nabla_\mu \varphi \text{ and } X = \varphi_\mu \varphi^\mu$$

$$L_4^{(2)} = \varphi^\mu \varphi_{\mu\nu} \varphi^{\nu\rho} \varphi_\rho$$

$$L_5^{(2)} = (\varphi^\rho \varphi_{\rho\sigma} \varphi^\sigma)^2$$

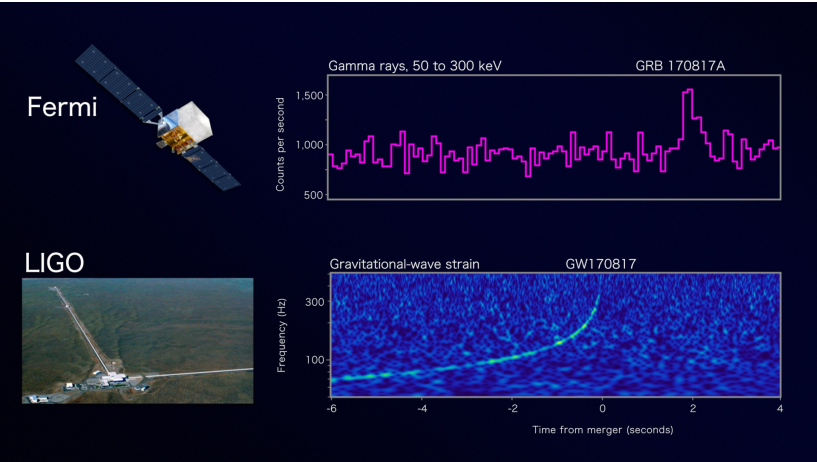
## DHOST theories

From now on,  $\varphi =$  dark energy



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# DHOST theories

From now on,  $\varphi$  = dark energy

$$c_{\text{grav}} = c_{\text{light}}$$

$$A_1 = 0$$



Degeneracy relations (no ghost)

$$A_2 = 0$$

$$A_4 = A_4[F_2, A_3]$$

$$A_5 = A_5[F_2, A_3]$$

# Breaking of Vainshtein mechanism inside matter

## Spherical object in expanding Universe

$$ds^2 = -[1 + 2\Phi(t, r)]dt^2 + a(t)^2[1 + 2\Psi(t, r)](dr^2 + r^2d\Omega^2)$$

Assumptions:

- Small field limit  $\Phi, \Psi \ll 1$ , and  $(\Phi'/\dot{\Phi})^2 \ll 1$
- Quasi-stationarity
- Vainshtein regime

## Modified Newton law

Kobayashi et al. '15

$$\phi' = \frac{G_N(t)M(t, r)}{r^2} + \frac{\Upsilon_1(t)G_N(t)}{4}M''(t, r)$$

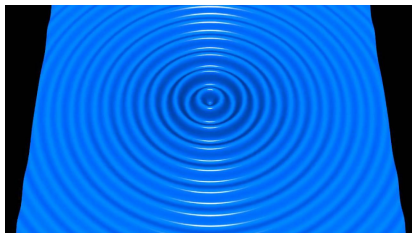
$$M = 4\pi a(t)^2 \int r^2 dr \rho(t, r)$$

$$G_N = G_N[F_2, A_3]$$

$$\Upsilon_1 = \Upsilon_1[F_2, A_3]$$

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# Spherical sound waves



## Hydrodynamics equation

$$\frac{\partial(\rho\vec{v})}{\partial t} + (\vec{v} \cdot \vec{\nabla})(\rho\vec{v}) = -\rho\vec{\nabla}\Phi - \vec{\nabla}P$$

$$\frac{\partial\rho}{\partial t} + \text{div}(\rho\vec{v}) = 0$$

$$P = P_0 + \delta P, \quad \rho = \rho_0 + \delta \rho, \quad \frac{\delta P}{\delta \rho} \simeq \left. \frac{\partial P}{\partial \rho} \right|_S$$

### Small wavelength perturbations

$$\frac{r}{\lambda} \gg 1, \quad \frac{r}{\lambda} \gg \frac{1}{\Upsilon_1} \frac{M}{\rho_0 r^3} \Rightarrow \delta \rho' \simeq \frac{\delta \rho}{\lambda} \text{ becomes dominant}$$

### Wave equation

$$\ddot{\delta \rho} - \left( \left. \frac{\partial P}{\partial \rho} \right|_S + 4\pi G_N \frac{\Upsilon_1}{4} \rho_0 r^2 \right) \delta \rho'' = 0$$

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Modified dispersion relation

Babichev and AL '18

$$c_{\text{DHOST}}^2 = \left. \frac{\partial P}{\partial \rho} \right|_S + 4\pi G_N \frac{\gamma_1}{4} \rho_0 r^2$$

$$c_{\text{GR}}^2 = \left. \frac{\partial P}{\partial \rho} \right|_S$$

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## Dust fluid (cosmology)

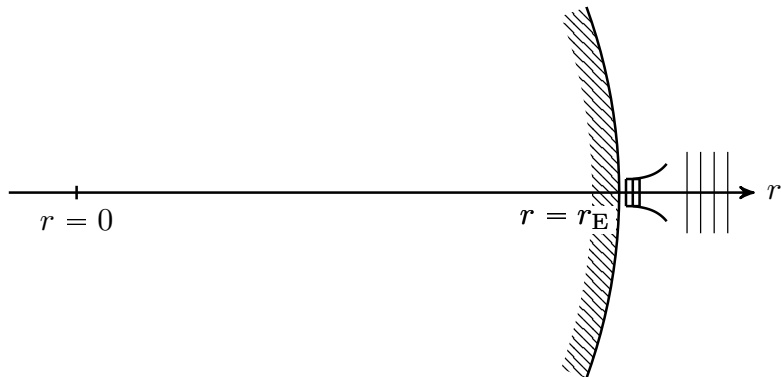


$$c_{\text{DHOST}}^2 = 4\pi G_{\text{N}} \frac{\Upsilon_1}{4} \rho_0 r^2$$

$\Upsilon_1 < 0 \Rightarrow$  Gradient instability, way worse than Jeans (tachyonic) instability

Stability  $\Rightarrow \Upsilon_1 > 0$

# Sound waves on Earth



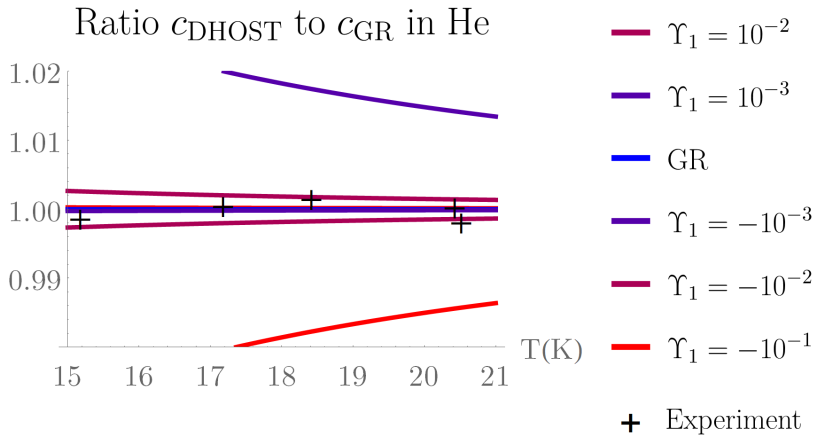
## Sound waves on Earth

Perfect gas

$$c_{\text{DHOST}}^2 = \frac{\gamma RT}{M} + 4\pi G_{\text{N}} \frac{\gamma_1 P_0 M}{4 RT} r_{\text{E}}^2$$

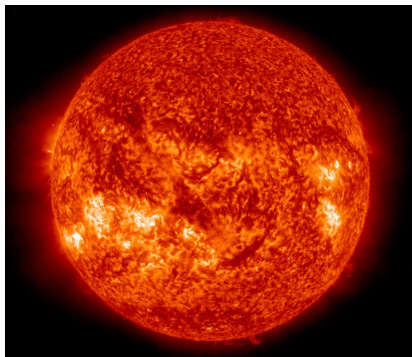
$$\gamma = \frac{7}{5}, \quad r_{\text{E}} \simeq 6.4 \times 10^6 \text{ m}, \quad M \simeq 4.0 \times 10^{-3} \text{ kg/mol for helium}$$

Strong temperature dependence, amplified at small temperature



$$|\Upsilon_1| \lesssim 10^{-2}$$

## Astrophysical constraints



→ See Ippocratis' talk



# Fundamental issue

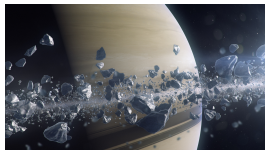
$$F_{\text{grav}} = -\phi' = -\frac{G_{\text{N}}M}{r^2} - \frac{\gamma_1 G_{\text{N}}}{4} M''$$

Gravitational force depends on matter description!



$$M'' \neq 0$$

$\neq$



$$M'' = 0$$

## Conclusions

- Constraint  $0 < \Upsilon_1 \lesssim 10^{-2}$ , can easily be refined
- Overall, DHOST theories for dark energy have a hard time
- Caveat: some models with different modified Newton law

Thank you!